

The hunt for axions

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Many theoretically well-motivated extensions of the Standard Model of particle physics predict the existence of the axion and further ultralight axion-like particles. They may constitute the mysterious dark matter in the universe and solve some puzzles in stellar and high-energy astrophysics. There are new, relatively small experiments around the globe, which started to hunt for these elusive particles and complement the accelerator based search for physics beyond the Standard Model.

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1. Introduction

In particle physics, we are in the comfortable situation that we have a theory – the Standard Model (SM) of particle physics – which describes all known particles and their interactions to a remarkable precision. On the other hand, astronomical observations tell us that only about fifteen percent of the matter in the universe is constituted by the known particles. It is one of the most urgent problems of fundamental physics to understand the nature of the remaining eighty-five percent of matter.

Theoretical particle physicists have been very imaginative to propose dark matter candidates spanning a huge range in mass and in strength of their interactions. Among those candidates, two particular ones stick out, both because of their appealing physics case and because of the variety of experimental probes: the neutralino – the lightest supersymmetric partner of the SM neutral gauge bosons and neutral Higgses – as a typical weakly interacting massive particle (WIMP), and the axion – the pseudo Nambu-Goldstone boson arising from the breaking of a global symmetry postulated in order to solve the strong CP problem – as a typical very weakly interacting slim (in the sense of ultralight) particle (WISP). We will devote our attention in this short review to the latter species (for more extended reviews, see Refs. [1, 2]).

2. Nambu-Goldstone bosons as natural WISP candidates

WISPs occur naturally in SM extensions featuring new global $U(1)$ symmetries which are spontaneously broken by a hidden Higgs mechanism at a symmetry breaking scale v_h much larger than the electroweak symmetry breaking scale, $v = 246$ GeV. In this case, the field $a(x)$ in the phase of the expansion of the hidden complex Higgs field about its vacuum expectation value (vev), $H_h(x) = \frac{1}{\sqrt{2}} (v_h + h_h(x)) e^{ia(x)/v_h}$, has a flat potential, $V(a) \equiv \text{const.}$, and is thus massless, while the field excitation $h_h(x)$ of the modulus has a large mass $m_h \propto v_h \gg v$. Moreover, at energies below the electroweak scale, the interactions of the so-called Nambu-Goldstone boson a with the SM particles – gauge bosons (gluons with field strength G and photons with field strength F) and fermions f (light quarks and leptons) – are suppressed by the inverse of the large symmetry breaking scale,

$$\mathcal{L} = \frac{1}{2} \partial_\mu a \partial^\mu a - \frac{\alpha_s}{8\pi} C_{ag} \frac{a}{f_a} G_{\mu\nu}^c \tilde{G}^{c,\mu\nu} - \frac{\alpha}{8\pi} C_{a\gamma} \frac{a}{f_a} F_{\mu\nu} \tilde{F}^{\mu\nu} + \frac{1}{2} \frac{C_{af}}{f_a} \partial_\mu a \bar{\psi}_f \gamma^\mu \gamma_5 \psi_f, \quad (2.1)$$

with model-dependent dimensionless coupling coefficients C_{aj} and a decay constant $f_a \simeq v_h$.

In models in which the current associated with the new $U(1)$ symmetry features an $SU(3)_C \times SU(3)_C \times U(1)$ chiral anomaly, $C_{ag} \neq 0$, the so-called strong CP problem is solved by non-perturbative QCD dynamics: the effective potential for a has then an absolute minimum at $a = 0$ and therefore the effective theta parameter – proportional to the vev of a – vanishes [3]. The corresponding Nambu-Goldstone boson is called the axion [4, 5]. Strictly speaking, it is a pseudo Nambu-Goldstone boson: it gets a small mass, $V(a) = \frac{1}{2} m_a^2 a^2 + \mathcal{O}(a^4)$, due to the same non-perturbative QCD effects,

$$m_a = \frac{m_\pi f_\pi}{f_a / C_{ag}} \frac{\sqrt{m_u m_d}}{m_u + m_d} \simeq 6 \mu\text{eV} \times \left(\frac{10^{12} \text{ GeV}}{f_a / C_{ag}} \right). \quad (2.2)$$

Here m_π and f_π are the neutral pion mass and decay constant, and m_u and m_d are the masses of the light quarks. There may be further axion-like particles (ALPs) which arise as Nambu-Goldstone bosons from the breaking of further well motivated global symmetries and which have no anomalous coupling to the pseudo-scalar gluon density, $C_{ag} = 0$, but a non-zero anomalous coupling to photons ($C_{a\gamma} \neq 0$) and/or leptons ($C_{al} \neq 0$) and/or light quarks ($C_{aq} \neq 0$) (see, for example, [6, 7]). In fact, string theory suggests a plenitude of such ALPs [8, 9, 10, 11].

3. Axion/ALP dark matter?

For large symmetry breaking scales, axions and ALPs have lifetimes much longer than the age of the universe and interact extremely weakly with the SM particles, qualifying them as dark matter candidates. Moreover, they are produced in the early universe via the vacuum realignment mechanism as a coherent state of many, extremely non-relativistic (and thus extremely cold) particles in the form of a classical, spatially coherent oscillating field [12, 13, 14]. Today's (time t_0) fraction of axion or ALP dark matter produced via the vacuum realignment mechanism is proportional to the average field amplitude squared, $\langle a^2 \rangle \equiv f_a^2 \langle \theta_a^2 \rangle$, at the time when the oscillations started, $t_{\text{osc}} \simeq (3/2) m_a^{-1}(t_{\text{osc}})$, [15]

$$R_a \equiv \frac{\rho_a}{\rho_{\text{DM}}}(t_0) \simeq 0.2 \sqrt{\frac{m_a(t_0)}{\text{eV}}} \sqrt{\frac{m_a(t_0)}{m_a(t_{\text{osc}})}} \left(\frac{f_a}{10^{11} \text{ GeV}} \right)^2 \langle \theta_a^2 \rangle. \quad (3.1)$$

Here, the indicated time-dependence of the mass arises from its temperature dependence, $m_a(t) \equiv m_a(T(t))$, taking into account possible plasma effects. From (3.1) it follows, that an appreciable fraction of light axion/ALP dark matter is only expected for sufficiently large symmetry breaking scale, $f_a \gtrsim 10^{10-12}$ GeV and correspondingly small couplings to SM particles (see Fig. 1).

4. Hints on axion/ALPs from astrophysics

4.1 Axion/ALP energy losses of stars?

Axions or ALPs could be produced in hot stellar plasmas and could thus transport energy out of stars, thereby possibly influencing stellar lifetimes and energy-loss rates. Stringent bounds on the axion or ALP couplings to photons and electrons in a wide mass range have been established from observations of stars in globular-clusters (GCs) [17]. These are relying on the fact that number counts of stars in particular branches of the color-magnitude diagram of GCs allow for detailed tests of stellar evolution. Recently, several authors have confronted new data with improved theoretical predictions and found hints for anomalous excessive energy losses.

A recent state-of-the-art analysis of Horizontal Branch (i.e. helium burning) stars (HBs) in a large sample of 39 Galactic GCs, exploiting modern stellar models, and taking into account the substantial dependence of the predictions on the He mass fraction, found a slight indication of additional losses which may be accounted by Primakoff-like axion/ALP emission in the Coulomb field of a charged particle, $\gamma + Ze \rightarrow Ze + a$, if the photon coupling is in the range [18]

$$g_{a\gamma} \equiv \frac{\alpha}{2\pi} \frac{C_{a\gamma}}{f_a} = 4.5_{-1.6}^{+1.2} \times 10^{-11} \text{ GeV}^{-1}, \quad \text{for } m_a \lesssim \text{few} \times 10 \text{ keV}. \quad (4.1)$$

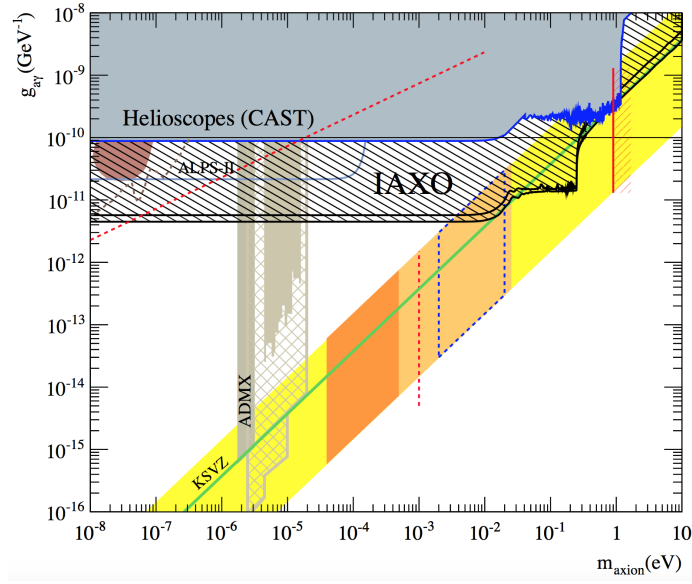


Figure 1: Current experimental constraints on the photon coupling of an axion or ALP [16]. Theoretically and astrophysically favored regions are shown for axions within the yellow model band (classical axion window in dark orange, mixed axion-WIMP DM in light orange, RG and WD cooling hint within the area surrounded by the dashed blue line) and for ALPs (brown dashed line for transparency hint, below red dashed diagonal line for ALP cold DM). Future prospects of ALPS II (above light blue line), IAXO (dashed black region), and ADMX (dashed brown region) are also shown.

Conservatively, the authors of this analysis determine also an upper bound, $g_{a\gamma} < 6.6 \times 10^{-11} \text{ GeV}^{-1}$, at 95 % confidence level (CL), which represents the strongest limit on $g_{a\gamma}$ for a wide mass range.

Clearly, the result (4.1) gives only a marginal hint for the existence of the QCD axion or an ALP. In fact, at around two sigma it is still compatible with the SM, $g_{a\gamma} = 0$. However, there are other mismatches between theory and observations which may also be seen as slight indications for the existence of a Nambu-Goldstone boson with such a coupling to the photon. One of them is the fact that the ratio of blue to red supergiants (SGs) is smaller than predicted by stellar evolution models. Moreover, the blue SGs appear to be less blue than expected [19]. This effect could also be explained by an axion/ALP with $g_{a\gamma} = \text{few} \times 10^{-11} \text{ GeV}^{-1}$ [20, 21].

Interestingly enough, Red Giants (RGs) in GCs also mildly prefer additional energy losses, this time, however, due to axion/ALP emission via bremsstrahlung of axions or ALPs, $e + Ze \rightarrow Ze + e + a$, pointing to an electron coupling in the range [22]

$$g_{ae} \equiv \frac{C_{ae} m_e}{f_a} = 1.8^{+0.6}_{-0.8} \times 10^{-13}, \quad \text{for } m_a \lesssim \text{few} \times 10 \text{ keV}. \quad (4.2)$$

Still, at $\sim 2\sigma$, the result is compatible with the SM, $g_{ae} < 4.3 \times 10^{-13}$ (95% CL).

Another astrophysical observable probing the electron coupling is the luminosity function of white dwarfs (WDs). Intriguingly, recent analyses, based on detailed WD cooling treatment and new data, find weak evidence that the WD luminosity function fits better with a new energy-loss channel that can be interpreted in terms of axion/ALP losses via bremsstrahlung in electron-ion or

electron-electron collisions, requiring an electron coupling in the range $g_{ae} = 1.0^{+0.2}_{-0.2} \times 10^{-13}$, for $m_a \lesssim \text{few} \times 10 \text{ keV}$ [23, 24]. This is consistent with the axion/ALP explanation of the apparent excessive energy loss of RGs, cf. Eq. (4.2). However, here again the evidence is still rather weak: at 2σ the WD luminosity function is consistent with the SM, $g_{ae} < 2.3 \times 10^{-13}$ (95% CL).

Finally, there is another hint for extra energy losses from the neutron star in Cassiopeia A (Cas A): its surface temperature measured over 10 years reveals an unusually fast cooling rate. This may be interpreted as a hint for extra cooling by axion/ALP due to nucleon bremsstrahlung, $N + N \rightarrow N + N + a$, requiring a coupling to the neutron of size [25]

$$g_{an} \equiv \frac{C_{an} m_n}{f_a} \sim 4 \times 10^{-10}, \quad \text{for } m_a \lesssim \text{few} \times 10 \text{ MeV}. \quad (4.3)$$

Still, this can also be seen as an approximate upper limit on this coupling. In fact, as recently pointed out, the more rapid cooling of the superfluid core in the neutron star may also arise from a phase transition of the neutron condensate into a multicomponent state [26].

Clearly, the evidence for each of the anomalous energy loss hints is very weak. But nevertheless, a unified picture seems to emerge: just one light, $m_a \lesssim \text{few} \times 10 \text{ keV}$, pseudo Nambu-Goldstone boson with $f_a \sim 10^8 \text{ GeV}$, $C_{a\gamma} \sim 1$, and $C_{ae} \sim C_{an} \sim 10^{-2}$, is required to explain at the same time the indications from the three quite different stellar energy loss channels.

Intriguingly, in LARGE volume string compactifications, in which the volume \mathcal{V} of the compactified extra dimensions is stabilized at an exponentially large value in units of the string size, the decay constant of closed string ALPs is generically much smaller than the Planck scale, $f_a \sim M_{\text{Pl}}/\sqrt{\mathcal{V}}$, and their matter coupling coefficients C_{af} are generically suppressed by a factor $\alpha \sim 10^{-2}$ in comparison to the photon coupling coefficient $C_{a\gamma}$, realizing the required properties [11].

4.2 Photon-ALP oscillations in astrophysical magnetic fields?

Photon-ALP oscillations in large-scale interstellar as well as in local magnetic fields associated with astrophysical objects provide another sensitive probe for ALPs complementary to Primakoff conversion in stars. However, this probe is only effective for large photon-ALP mixing and is therefore confined to very low masses,

$$|m_a^2 - \omega_{\text{pl}}^2| \ll 2g_{a\gamma} B E \simeq 2.5 \times \text{neV}^2 \left(\frac{g_{a\gamma}}{10^{-11} \text{ GeV}^{-1}} \right) \left(\frac{B}{\mu\text{G}} \right) \left(\frac{E}{\text{GeV}} \right), \quad (4.4)$$

where $\omega_{\text{pl}} \simeq 4 \times 10^{11} \text{ eV} \sqrt{n_e/\text{cm}^3}$ is the plasma frequency in terms of the electron density in the medium, while B is the magnetic field strength and E is the photon/axion/ALP energy.

In a galactic core-collapse supernova (SN), ALPs would be emitted via the Primakoff process, and oscillate into gamma rays in the magnetic field of the Milky Way. The lack of a gamma-ray signal in the GRS instrument of the SMM satellite in coincidence with the observation of the neutrinos emitted from SN1987A therefore provides a strong bound on their coupling to photons, cf. [27, 28]. Recently, this bound has been revisited and the underlying physics has been brought to the current state-of-the-art, as far as modelling of the supernova and the Milky-Way magnetic field are concerned, resulting in the limit [29]

$$g_{a\gamma} < 5.3 \times 10^{-12} \text{ GeV}^{-1}, \quad \text{for } m_a \lesssim 4.4 \times 10^{-10} \text{ eV}. \quad (4.5)$$

Therefore, the mass window for ALPs explaining the HB energy loss shrinks to $\text{neV} \lesssim m_a \lesssim \text{few} \times 10 \text{ keV}$.

Gamma-ray spectra from distant active galactic nuclei (AGN) should show an energy and redshift-dependent exponential attenuation, $\exp(-\tau(E, z))$, due to e^+e^- pair production off the extragalactic background light (EBL) – the stellar and dust-reprocessed light accumulated during the cosmological evolution following the era of re-ionization. However, a number of authors [30, 31, 32] have noted that the observed spectra seem to point to an anomalous transparency of the universe for gamma-rays at large optical depth, $\tau \gtrsim 2$, although with evidence below two sigma, if one takes into account systematic effects (EBL spectrum, intrinsic source spectra, etc.) [33]. This may be explained by photon \leftrightarrow ALP oscillations: the conversion of gamma rays into ALPs in the magnetic fields around AGNs or in the intergalactic medium, followed by their unimpeded travel towards our galaxy and the consequent reconversion into photons in the (inter)galactic magnetic fields, require an ALP with [34, 35, 36, 37, 38, 39, 40]

$$g_{a\gamma} \gtrsim 10^{-(11-12)} \text{ GeV}^{-1}, \quad \text{for } m_a \lesssim 10^{-7} \text{ eV}. \quad (4.6)$$

Intriguingly, a similar value of the photon coupling is required also for a possible explanation of the rapidly varying very high energy ($E > 50 \text{ GeV}$) emission from the flat spectrum radio quasar PKS 1222+216 which represents a challenge for standard blazar scenarios: in the latter one is forced to invoke the existence of a very compact emitting region at a large distance from the jet base, in order to avoid absorption of gamma rays in the dense ultraviolet radiation field of the broad line region. In ref. [41] it was shown that one can also use a standard blazar model for PKS 1222+216 where gamma rays are produced close to the central engine, if one assumes that inside the source photons can oscillate into ALPs. Moreover, the required photon coupling overlaps with the preferred region from the anomalous energy losses of helium burning stars in globular clusters.

Finally, it was found that observed soft X-ray excesses in many galaxy clusters may be explained by the conversion of a cosmic ALP background (CAB) radiation, corresponding to an effective number ΔN_{eff} of extra neutrinos, into photons in the cluster magnetic fields [42, 43, 44]. This explanation requires that the CAB spectrum is peaked in the soft X-ray region and that the ALP coupling and mass satisfy $g_{a\gamma} \gtrsim (1-2) \times 10^{-13} \text{ GeV}^{-1} \sqrt{0.5/\Delta N_{\text{eff}}}$, for $m_a \lesssim 10^{-12} \text{ eV}$.

5. The experimental hunt for the axion and ALPs

We have seen that there is a strong physics case for the axion and other ALPs arising as Nambu-Goldstone bosons from the breaking of a symmetry at a scale $f_a \sim 10^{8 \div 12} \text{ GeV}$. Fortunately, such WISPs are in reach of a number of experiments which are presently carried out or being set up.

5.1 Light-shining-through-a-wall searches

Light-Shining-Through-Wall (LSW) experiments aim both for production and detection of axions and ALPs in the laboratory. This is done by sending laser photons along a strong magnetic field, allowing for their conversion into axions or ALPs, towards a blocking wall, behind of which the latter may then reconvert, again in a strong magnetic field, into photons, the latter being susceptible to detection (for a review, see [45]). The Any Light Particle Search (ALPS) experiment

at DESY and OSQAR at CERN share currently the best sensitivity of LSW experiments [46, 47]. ALPS II [48] proposes to use 10+10 straightened HERA magnets [49], a high-power laser system developed and the pioneering realization of an optical regeneration cavity [50, 51] developed by the Albert-Einstein Institute, and a superconducting low-background detector. It aims at crucially testing the ALP explanation of the excessive helium burning star energy loss and of the anomalous cosmic gamma ray transparency, cf. Fig. 1.

5.2 Helioscope searches

Helioscopes aim at detecting solar axions and ALPs produced by their conversion into photons inside of a strong magnet pointing towards the Sun [52]. The CERN Axion Solar Telescope (CAST), employing an LHC dipole test magnet, currently sets the best helioscope limit [53, 54]. A proposed next-generation axion helioscope, dubbed the International Axion Observatory (IAXO), envisions a dedicated superconducting toroidal magnet with much bigger aperture than CAST, a detection system consisting of large X-ray telescopes coupled to ultra-low background X-ray detectors, and a large, robust tracking system [55]. It aims at the sensitivity shown in Fig. 1. It will crucially test the axion explanation of the energy losses of helium burning stars, red giants, white dwarfs and the neutron star in Cas A, and of the ALP explanation of the anomalous cosmic gamma ray transparency. Furthermore, it will also scratch the parameter region required to explain the soft X-ray excess from galaxy clusters.

5.3 Direct axion/ALP dark matter searches

5.3.1 Haloscopes

Haloscopes directly search for galactic halo dark matter axions and ALPs in the laboratory via their coupling to the photon. Currently, the most sensitive ones exploit electromagnetic cavities placed in a strong magnet [52]. They aim for the detection of electromagnetic power arising from the conversion of dark matter axions or ALPs into real photons, with frequency $\nu = m_a/(2\pi) = 0.24 \text{ GHz} \times (m_a/\mu\text{eV})$. The best sensitivity is reached on resonance, the power output then being proportional to the quality factor of the cavity. The Axion Dark Matter eXperiment (ADMX) has indeed reached recently the sensitivity to probe axion dark matter [56] (see Fig. 1). Further microwave cavity based haloscope opportunities in complementary mass ranges may arise from recycling available microwave cavities and magnets at accelerator laboratories [57, 58].

Other new concepts for haloscopes are also being investigated [59]. A microwave Fabry-Perot resonator in a spatially varying magnetic field may be exploited to search for axion/ALP dark matter with masses above $40 \mu\text{eV}$ [60]. Converted photons from axion/ALP dark matter could be focused in a manner similar to a dish antenna, allowing for broad-band searches [61].

5.3.2 Magnetic resonance searches

Axion dark matter produced via the misalignment mechanism gives oscillating electric dipole moments (EDMs) to all nucleons. These EDMs cause the precession of nuclear spins in a nucleon spin polarized sample in the presence of an electric field. The resulting transverse magnetization can be searched for by exploiting magnetic-resonance (MR) techniques [62]. The aim of the corresponding Cosmic Axion Spin Precession Experiment (CASPEr) in Mainz is to probe axion dark

matter in the anthropic window, corresponding to GUT to Planck scale symmetry-breaking scales $f_a \sim 10^{15-18}$ GeV, complementary to the classic axion window probed by ADMX.

The axion/ALP nucleon coupling C_{aN} will also lead to a spin precession about the axion/ALP DM wind, even without the presence of an electric field. Therefore, CASPER can also be exploited to search for the magnetization due to this effect [63]. Unfortunately, the projected sensitivity does not reach the axion prediction in this case.

Finally, the axion/ALP electron coupling C_{ae} will also lead to a spin precession about the axion/ALP DM wind [64, 65]. The QUAX (QUAerere AXions) experiment in preparation in Italy by INFN aims at exploiting MR inside a magnetized material [66]. Because of the higher Larmor frequency of the electron, it extends the sensitivity to higher masses.

6. Summary

There is a strong physics case for the axion and ALPs. They occur naturally in many theoretically appealing ultraviolet completions of the SM. They are dark matter candidates and can explain the hints for an excessive energy loss of stars and for an anomalous transparency of the universe for TeV photons. A significant portion of their parameter space will be tackled in this decade by experiments. Stay tuned!

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